

# Black hole Magnetospheres and The Orthogonal GRB model

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# Our team at the Academy

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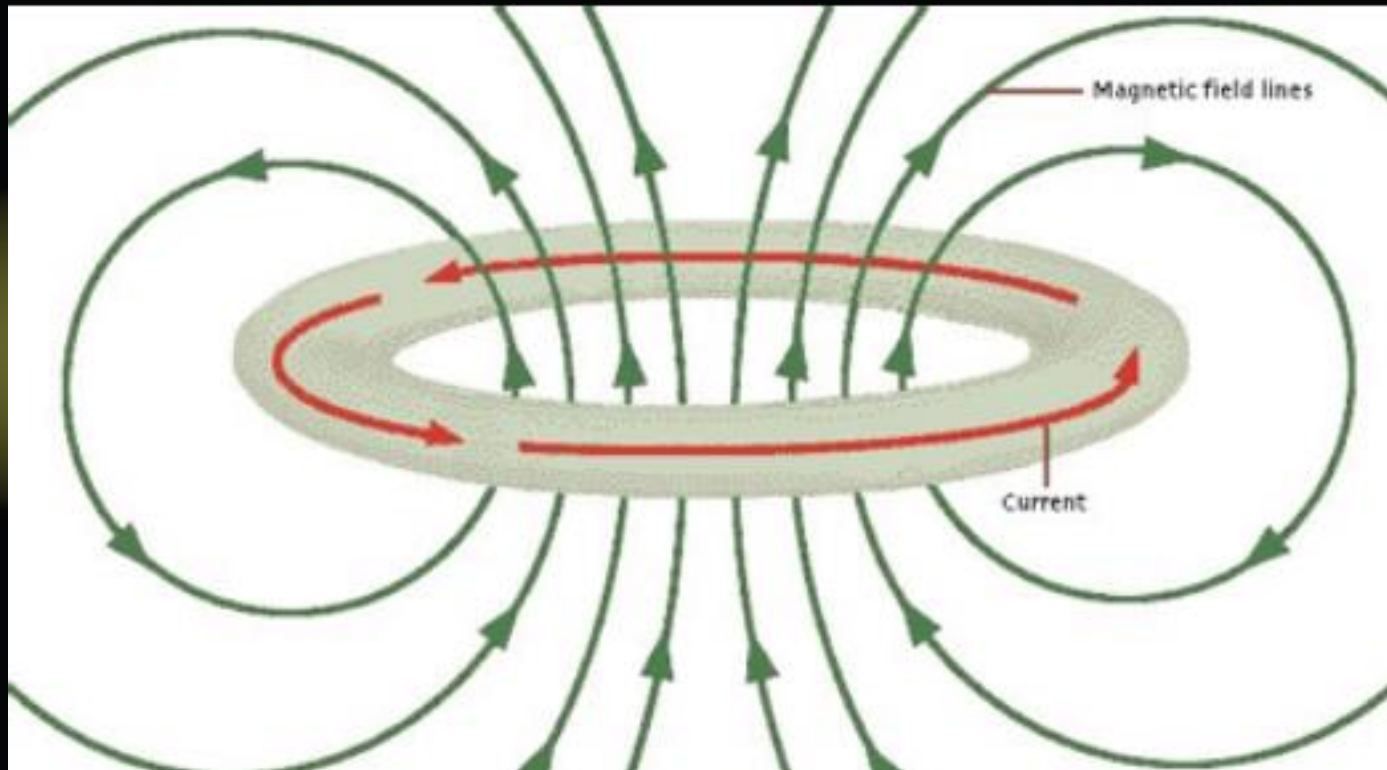


# Our aim is to understand:

- The origin of cosmic magnetic fields
- Cosmic Battery
- The Blandford & Znajek mechanism
- What is the structure of the black hole magnetosphere
- The analogy with pulsars



# The Cosmic Battery



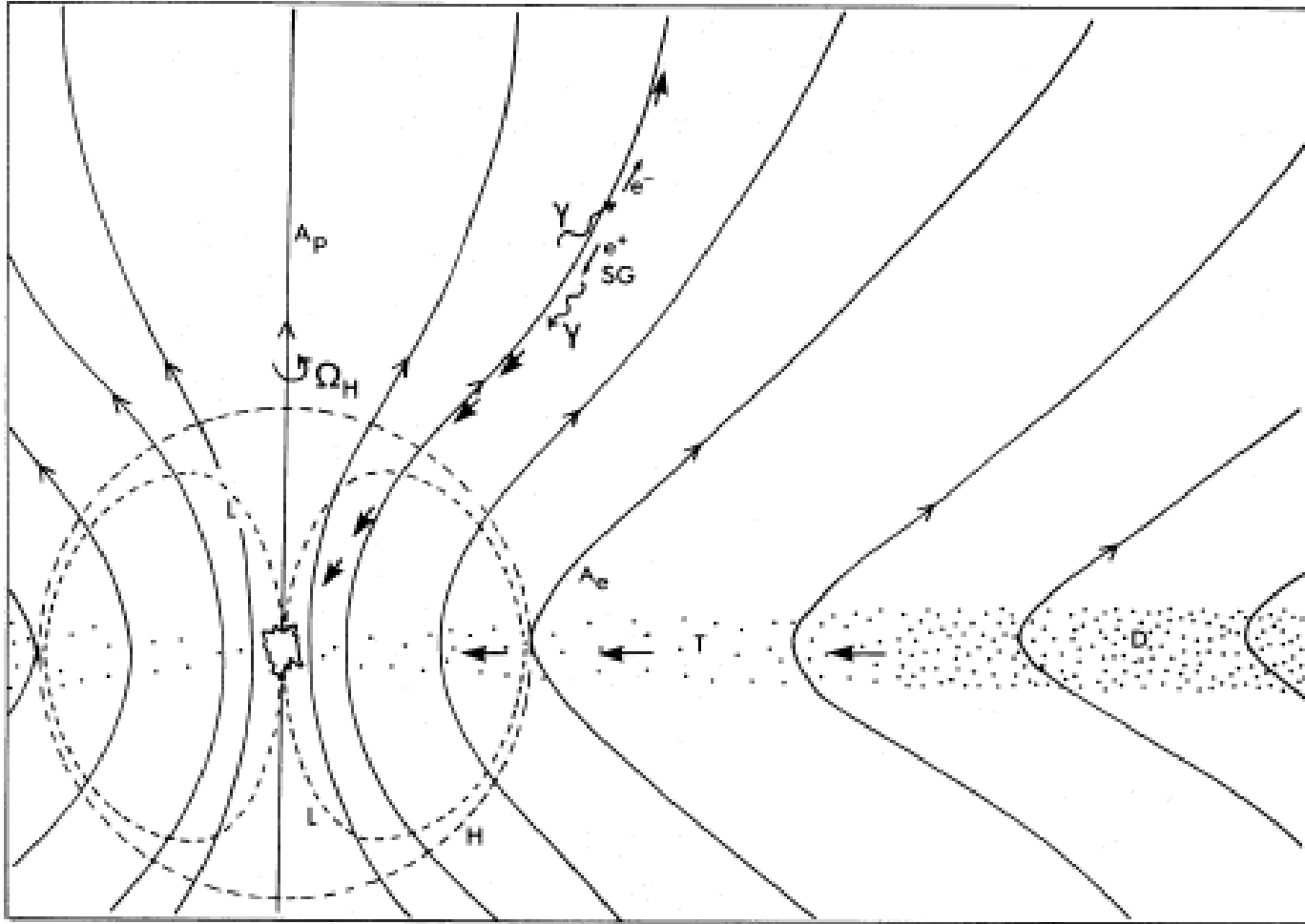
# The Cosmic Battery

# Blandford & Znajek mechanism (1977)

$$\mathcal{E}_{EM} \propto \omega(\Omega_{\text{BH}} - \omega)\Psi_m^2 \sim \Omega_{\text{BH}}^2 \Psi_m^2$$

- Rotating black hole
- Magnetic field supported by external currents
- Pair production & force free magnetosphere
- Intense flux of electromagnetic energy

# Blandford & Znajek mechanism (1977)



# MacDonald & Thorne (1982)

Things look  
familiar

All the insight from  
pulsar electrodynamics  
and other astrophysical  
problems



# GR Pulsar equation

$$\begin{aligned} ds^2 &= g_{\mu\nu} dx^\mu dx^\nu \\ &= -\alpha^2 dt^2 + \frac{A \sin^2 \theta}{\Sigma} (d\phi - \Omega dt)^2 + \frac{\Sigma}{\Delta} dr^2 + \Sigma d\theta^2 \end{aligned}$$

Kerr metric

$$\tilde{\nabla} \cdot \tilde{B} = 0$$

$$\tilde{\nabla} \cdot \tilde{E} = 4\pi \rho_e$$

$$\tilde{\nabla} \times (\alpha \tilde{B}) = 4\pi \alpha \tilde{J}$$

$$\nabla \times (\alpha \tilde{E}) = 0 .$$

# GR Pulsar equation

$$\tilde{E} \cdot \tilde{B} = 0$$

$$\tilde{B}(r, \theta) = \frac{1}{\sqrt{A} \sin \theta} \left\{ \Psi_{,\theta}, -\sqrt{\Delta} \Psi_{,r}, \frac{2I\sqrt{\Sigma}}{\alpha} \right\}$$

$$\tilde{E}(r, \theta) = \frac{\Omega - \omega}{\alpha\sqrt{\Sigma}} \left\{ \sqrt{\Delta} \Psi_{,r}, \Psi_{,\theta}, 0 \right\} .$$

Force free

$$\rho_e \tilde{E} + \tilde{J} \times \tilde{B} = 0$$

# The ‘Huge’ GR Pulsar equation

$$\begin{aligned}
 & \left\{ \Psi_{,rr} + \frac{1}{\Delta} \Psi_{,\theta\theta} + \Psi_{,r} \left( \frac{A_{,r}}{A} - \frac{\Sigma_{,r}}{\Sigma} \right) - \frac{\Psi_{,\theta} \cos \theta}{\Delta \sin \theta} \right\} \cdot \left[ 1 - \frac{\omega^2 A \sin^2 \theta}{\Sigma} + \frac{4M\alpha\omega r \sin^2 \theta}{\Sigma} - \frac{2Mr}{\Sigma} \right] \\
 & - \left( \frac{A_{,r}}{A} - \frac{\Sigma_{,r}}{\Sigma} \right) \Psi_{,r} - \left( 2 \frac{\cos \theta}{\sin \theta} - \frac{A_{,\theta}}{A} + \frac{\Sigma_{,\theta}}{\Sigma} \right) (\omega^2 A \sin^2 \theta - 4M\alpha\omega r \sin^2 \theta + 2Mr) \frac{\Psi_{,\theta}}{\Delta \Sigma} \\
 & + \frac{2Mr}{\Sigma} \left( \frac{A_{,r}}{A} - \frac{1}{r} \right) \Psi_{,r} + \frac{4\omega M\alpha r \sin^2 \theta}{\Sigma} \left\{ \Psi_{,r} \left( \frac{1}{r} - \frac{A_{,r}}{A} \right) - \frac{\Psi_{,\theta} A_{,\theta}}{\Delta A} \right\} \\
 & - \frac{\omega' \sin^2 \theta}{\Sigma} (\omega A - 2\alpha Mr) \left( \Psi_{,r}^2 + \frac{1}{\Delta} \Psi_{,\theta}^2 \right) = -\frac{4\Sigma}{\Delta} II'
 \end{aligned}$$

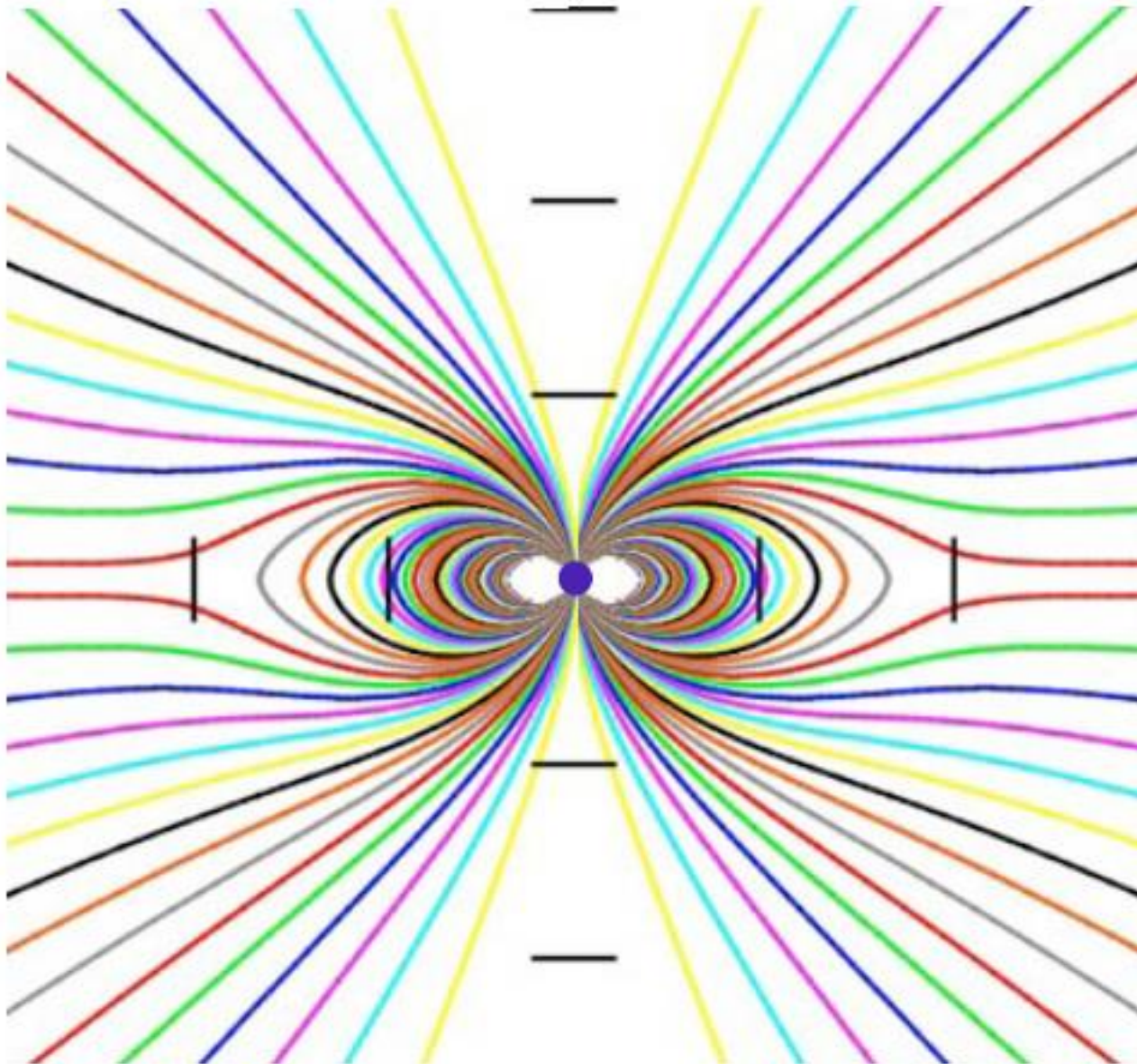
# The “small” Pulsar equation

$$\left( \Psi_{,rr} + \frac{1}{r^2} \Psi_{,\theta\theta} + \frac{2\Psi_{,r}}{r} - \frac{1 \cos \theta}{r^2 \sin \theta} \Psi_{,\theta} \right) \cdot [1 - \omega^2 r^2 \sin^2 \theta]$$
$$- \frac{2\Psi_{,r}}{r} - 2\omega^2 \cos \theta \sin \theta \Psi_{,\theta} - \omega\omega' r^2 \sin^2 \theta \left( \Psi_{,r}^2 + \frac{1}{r^2} \Psi_{,\theta}^2 \right) = -4II'$$

Pulsar Light cylinder:  $r \sin \theta = c / \omega$

The electric current  $I(\Psi)$  must be determined self-consistently

# The Pulsar Problem



Contopoulos, Kazanas, Fendt (1999)

# GR Pulsar equation

$$1 - \frac{\omega^2 A \sin^2 \theta}{\Sigma} + \frac{4M\alpha\omega r \sin^2 \theta}{\Sigma} - \frac{2Mr}{\Sigma} = 0$$

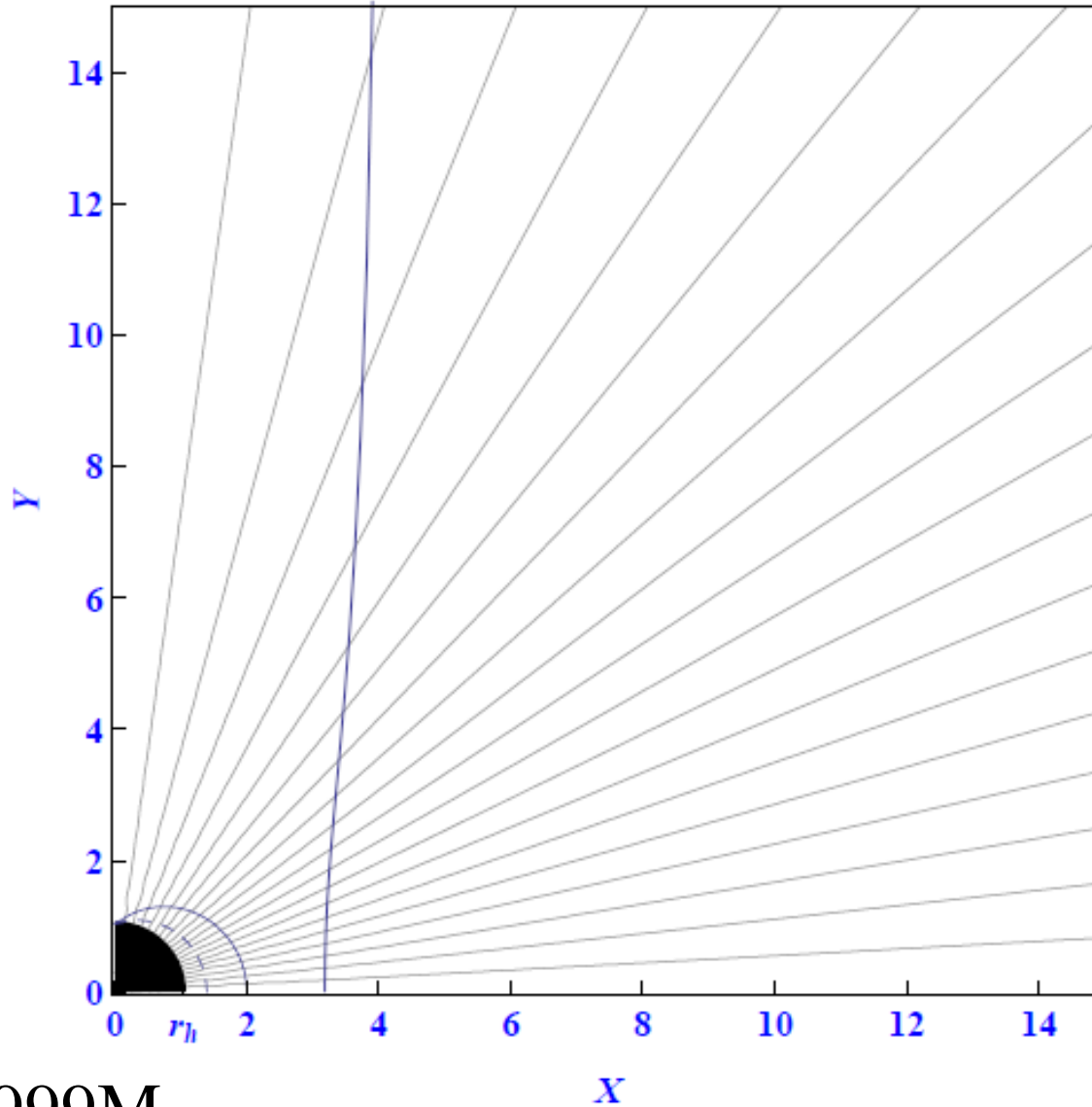
Two light surfaces one inside the ergosphere and one similar to the light cylinder

Two free functions to find self consistently :

$$I(\Psi) \quad \text{and} \quad \omega(\Psi)$$

The current and the angular velocity of the magnetic field

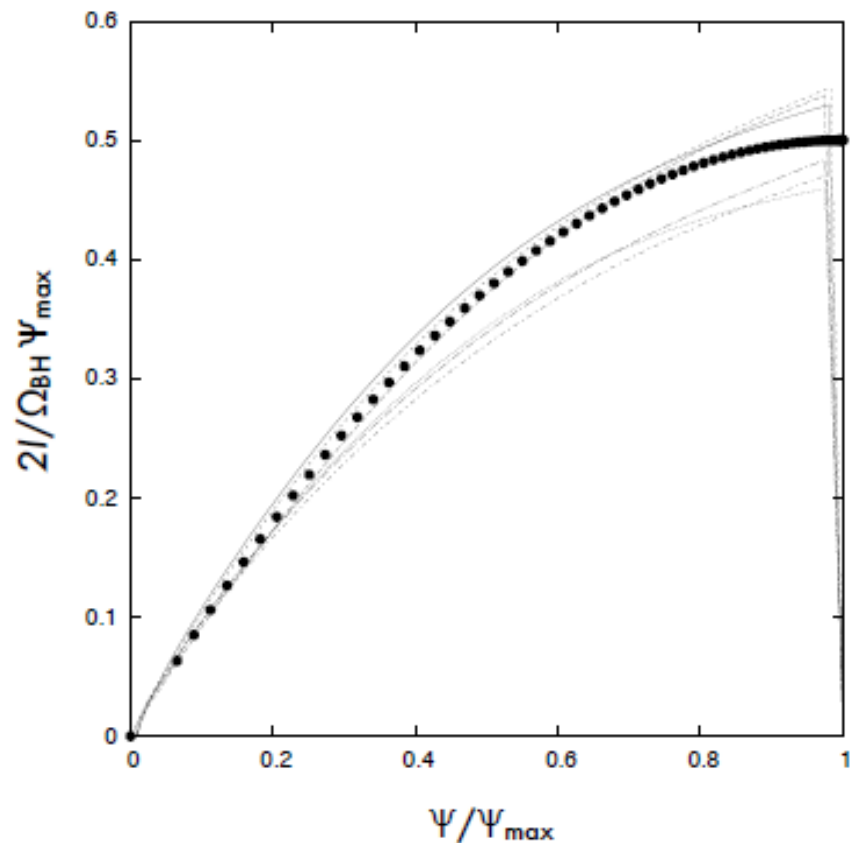
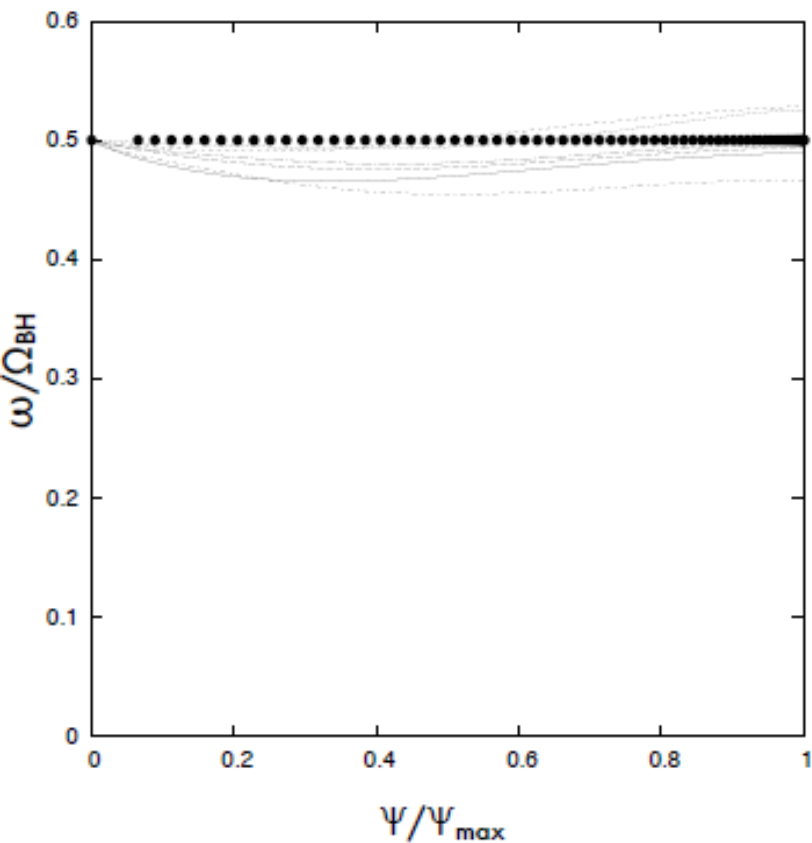
# Black hole monopole solution



$a=0.999M$

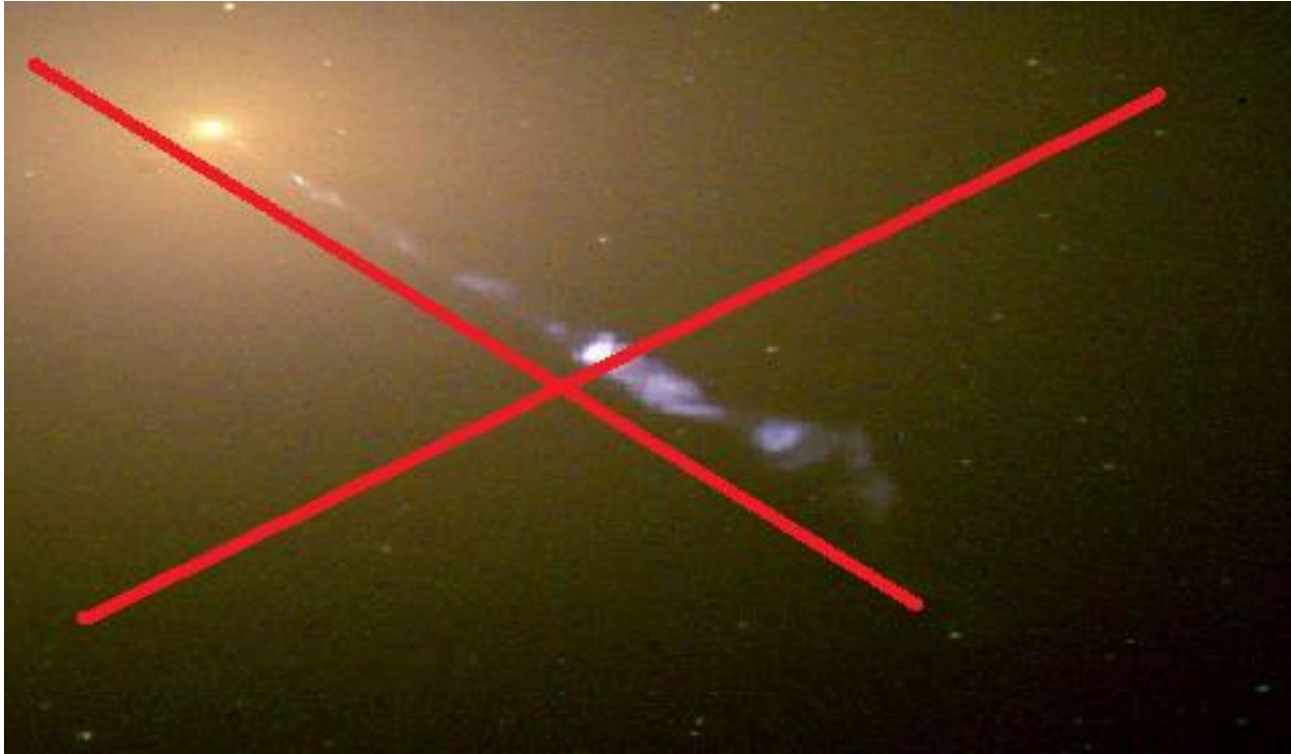
# GR Pulsar equation

$$\alpha=0.7-1 M, \omega \sim 0.5 \Omega_{\text{BH}}$$

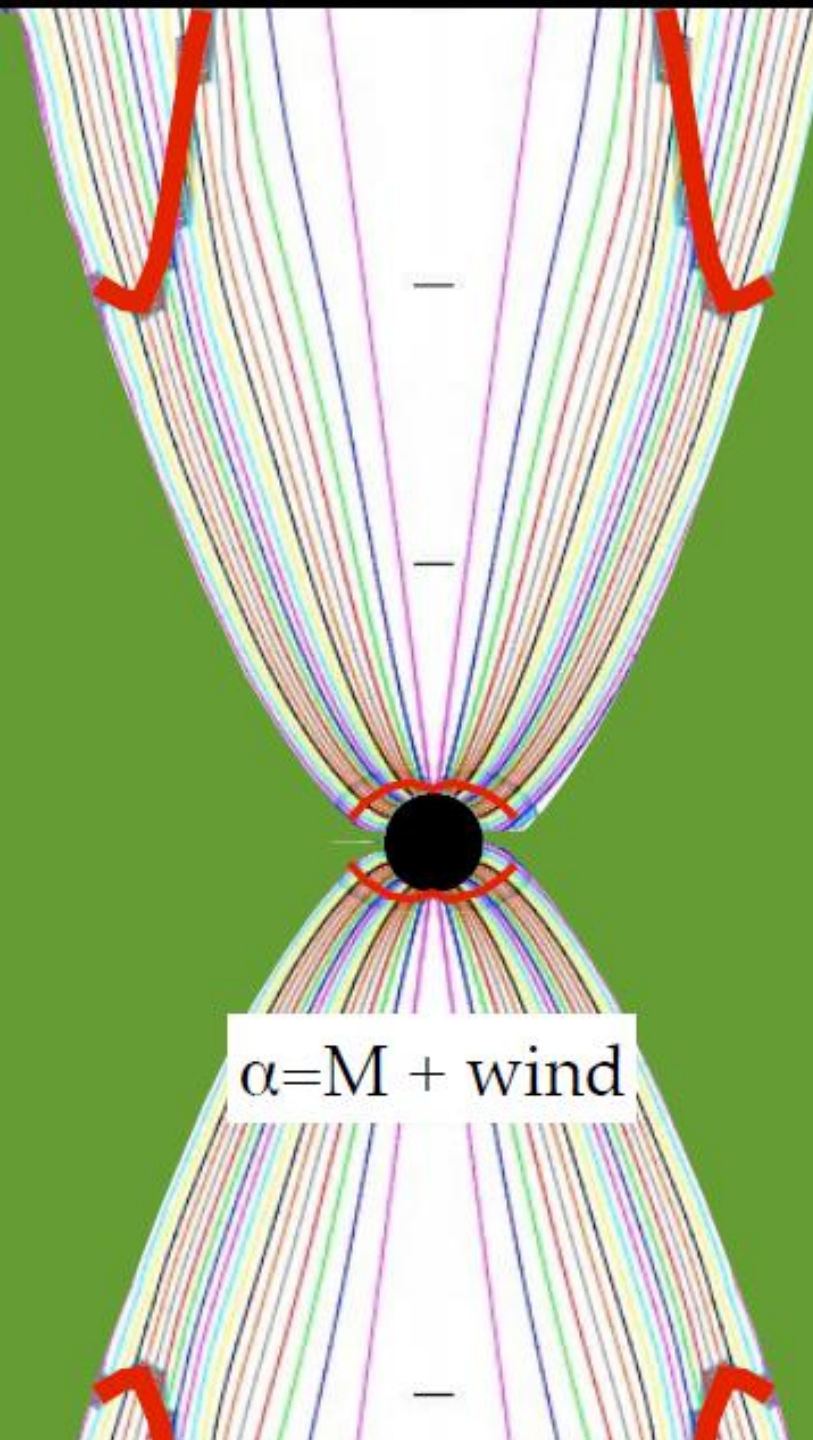




No jet found in these solutions

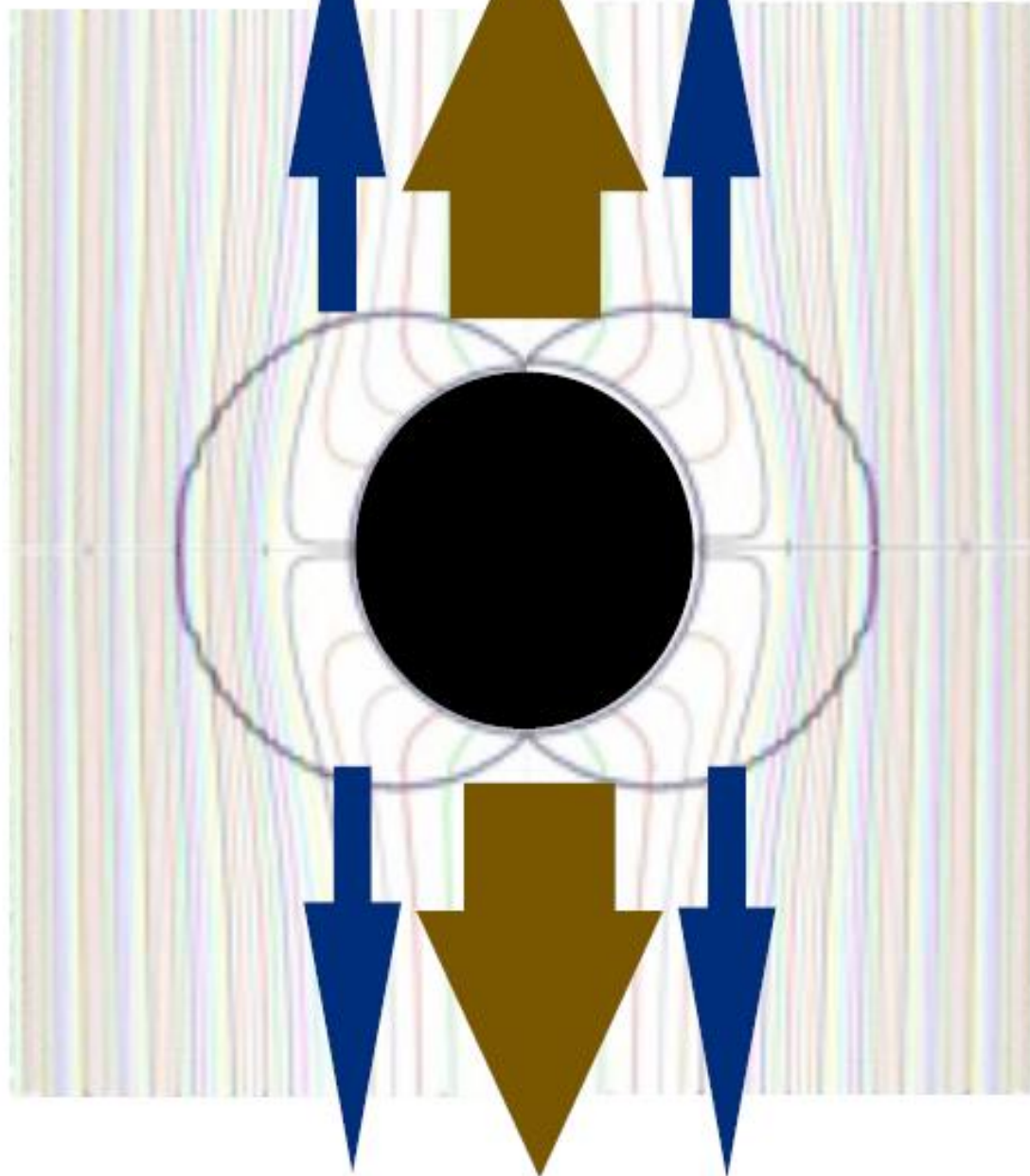


Something is needed to collimate  
the jet

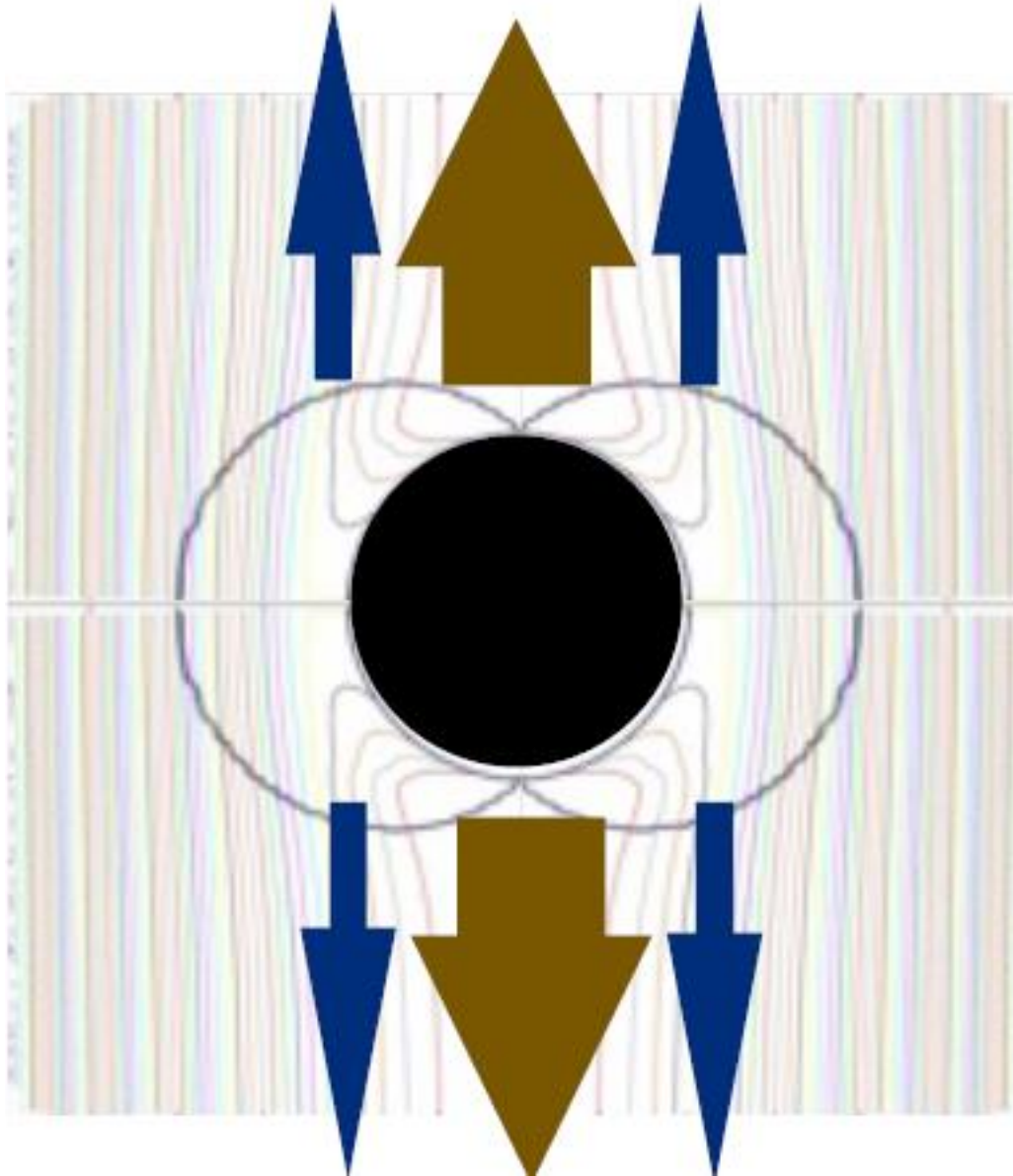


$\alpha = M + \text{wind}$

Not real jets



Not real jets



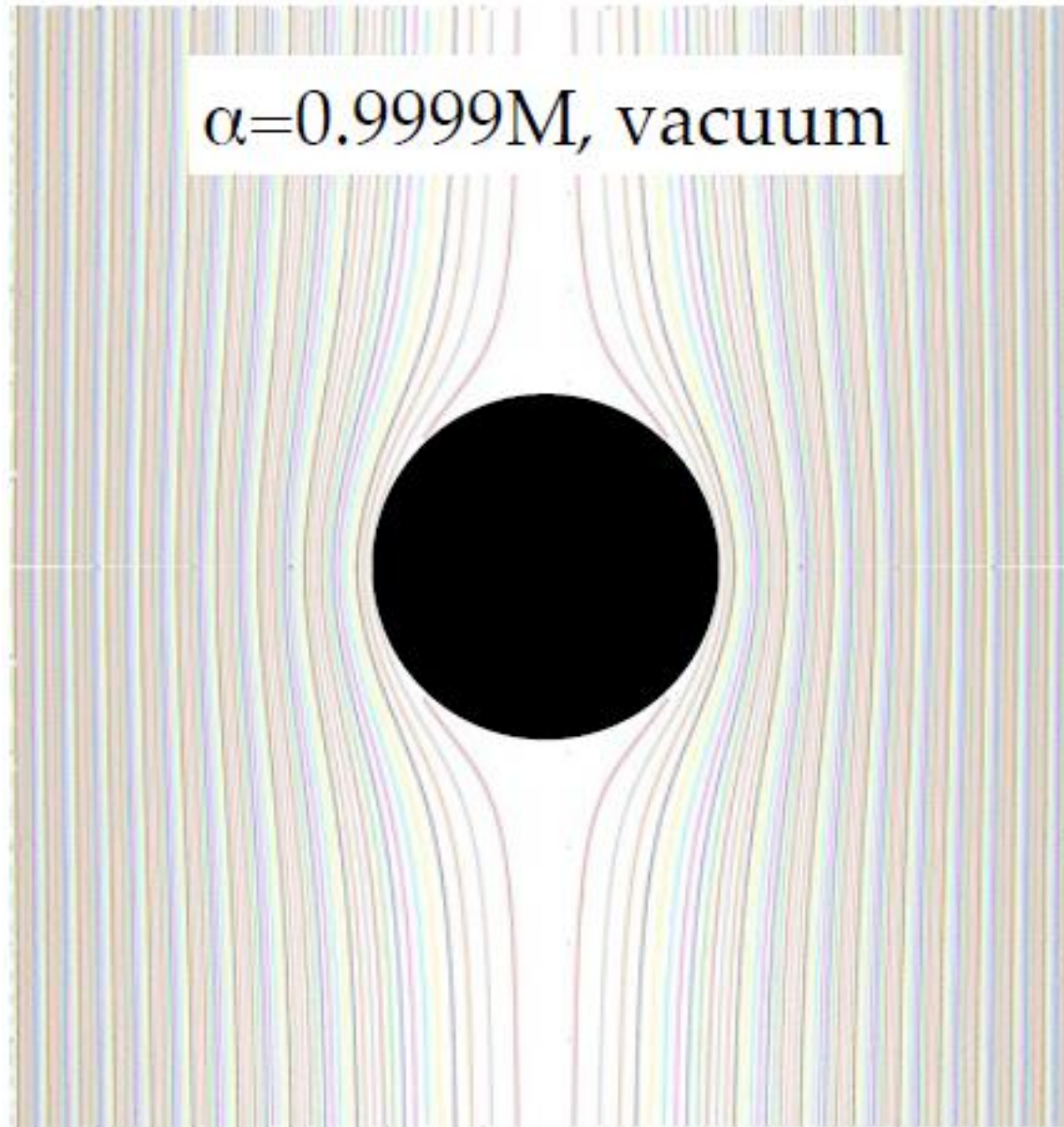
# Vacuum Solutions

$$\tilde{\nabla} \times (\alpha \tilde{B}) |^{\phi} = 0$$

It is the same  
if you put in  
the equation

$I=0, \omega=\Omega$  everywhere

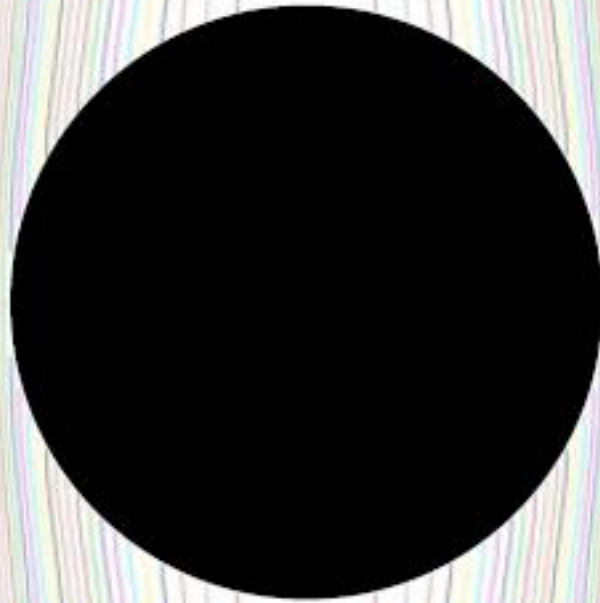
# Vacuum Solutions





# Vacuum Solutions

$\alpha=0.9M$ , vacuum

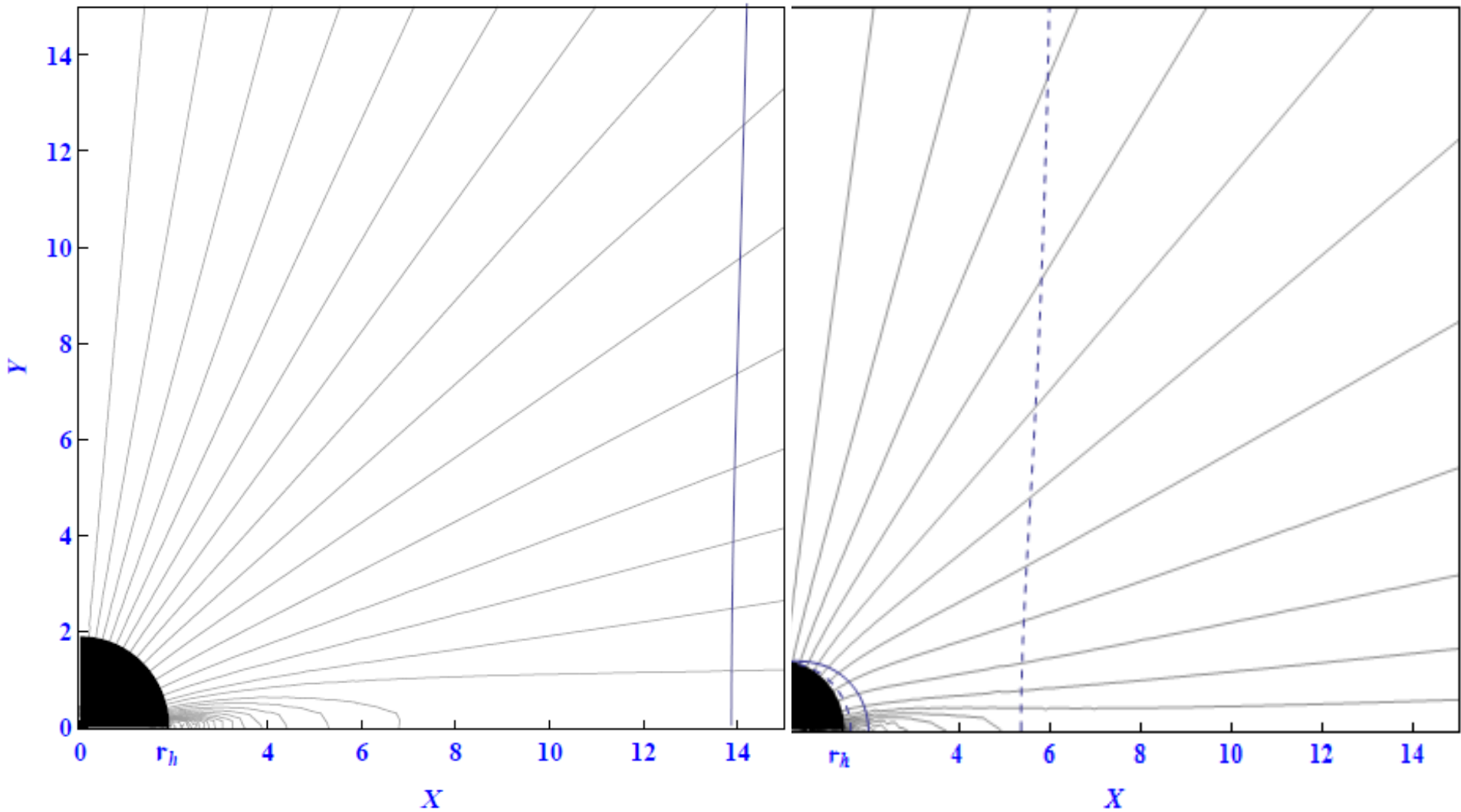


# The analogy with the pulsar

- When the accretion disc is starting to disperse a magnetospheric charge is there rotating
- This charge can support a dipole magnetic field distorted by the rapid rotation
- We have a magnetosphere with closed and open field lines
- Neither system forms relativistic jets on its own



# The magnetic black hole



$a=0.5$

Nathanail et.al in preparation

$a=0.9$

# The orthogonal GRB model

$$\dot{E} \sim M r_{bh}^2 \omega \dot{\omega}$$

$$\dot{E}_{BZ} \sim \Psi^2 \omega^2 \sim \omega^2 \sim e^{-t/\tau}$$

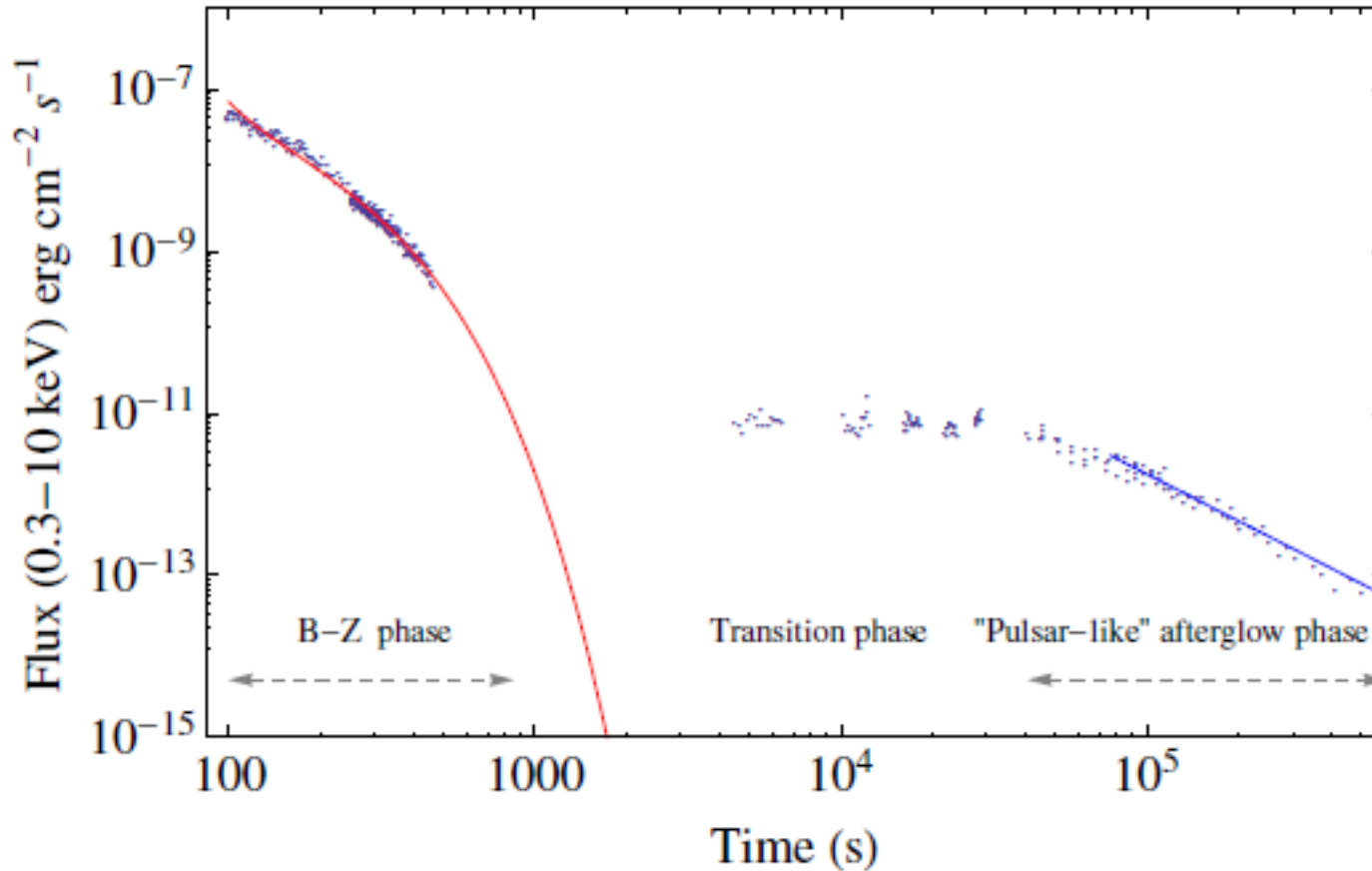
$$\dot{E}_{\text{puls.}} \sim \Psi_{\text{open}}^2 \omega^2 \sim \omega^8 \sim t^{-4/3}$$

# The orthogonal GRB model

- As in pulsars, high energy radiation is generated in the equatorial current sheet
- Our model is ‘orthogonal’ to the standard GRB model where all the action takes place along a relativistic jet emitted along the rotation and magnetic axis

# The orthogonal GRB model

GRB060614A



$$\dot{E}_o \tau_{\text{BZ}} = 10^{55} M_{10} \text{ erg}$$

$$\tau_{\text{BZ}} \sim 10 - 100 \text{ sec}$$

$$\dot{E}_o \sim 10^{53} \text{ erg/sec}$$

$$B_o \sim 10^{16} \text{ G}$$

# The end

# Thank you



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$$\begin{aligned}
& \left\{ \Psi_{,rr} + \frac{1}{\Delta} \Psi_{,\theta\theta} + \Psi_{,r} \left( \frac{A_{,r}}{A} - \frac{\Sigma_{,r}}{\Sigma} \right) - \frac{\Psi_{,\theta} \cos \theta}{\Delta \sin \theta} \right\} \cdot \left[ 1 - \frac{\omega^2 A \sin^2 \theta}{\Sigma} + \frac{4M\alpha\omega r \sin^2 \theta}{\Sigma} - \frac{2Mr}{\Sigma} \right] \\
& - \left( \frac{A_{,r}}{A} - \frac{\Sigma_{,r}}{\Sigma} \right) \Psi_{,r} - \left( 2 \frac{\cos \theta}{\sin \theta} - \frac{A_{,\theta}}{A} + \frac{\Sigma_{,\theta}}{\Sigma} \right) (\omega^2 A \sin^2 \theta - 4M\alpha\omega r \sin^2 \theta + 2Mr) \frac{\Psi_{,\theta}}{\Delta \Sigma} \\
& + \frac{2Mr}{\Sigma} \left( \frac{A_{,r}}{A} - \frac{1}{r} \right) \Psi_{,r} + \frac{4\omega M\alpha r \sin^2 \theta}{\Sigma} \left\{ \Psi_{,r} \left( \frac{1}{r} - \frac{A_{,r}}{A} \right) - \frac{\Psi_{,\theta} A_{,\theta}}{\Delta A} \right\} \\
& - \frac{\omega' \sin^2 \theta}{\Sigma} (\omega A - 2\alpha Mr) \left( \Psi_{,r}^2 + \frac{1}{\Delta} \Psi_{,\theta}^2 \right) = -\frac{4\Sigma}{\Delta} II'
\end{aligned}$$





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